Lecture IV: Perturbations

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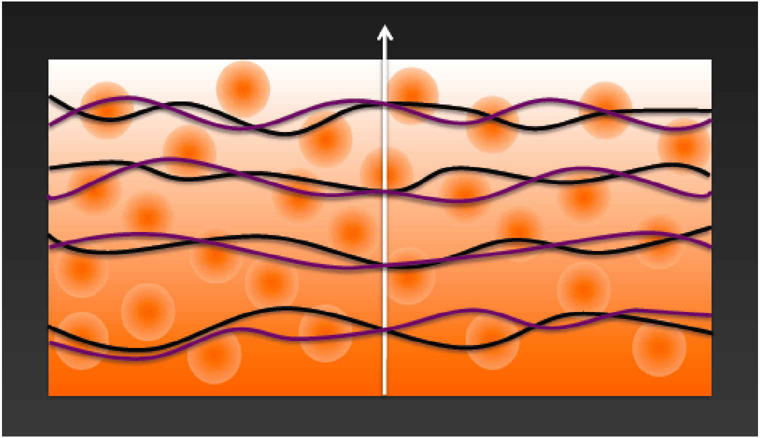
- first step for the C_l's, which depend on particular sources $S_i(k, \tau)$ integrated over time.

In this module, everything is assumed to be linear.

Gauge



- One gauge = one way to slice the space-time in equal-time hypersurfaces.
- Any time slicing such that quantities vary perturbatively on each slice is valid.
- Fixing the gauge = imposing a restriction (e.g. on number of non-zero metric perturbations) such that time slicing is fixed.
- Any choice of gauge leads to same observable quantities.

Gauge

$\delta g_{\mu\nu}$: 10 d.o.f., 4 scalars, 4 vectors, 2 tensors.

For scalar sector:

- imposing zero non-diagonal terms fixes the gauge.
  $ds^2 = -(1+2\psi)dt^2 + (1-2\phi)a^2 d\vec{l}^{\,2}$.
  Newtonian gauge.
- imposing zero terms in δg_{00} and δg_{0i} leaves also 2 d.o.f (h', η), but does not fix the gauge.
  But extra condition $\theta_i(k, \tau_{\rm ini}) = 0$ does.
  Synchronous gauge comoving with species i.

For a pressureless component, $\theta_i' = -\frac{a'}{a}\theta_i$. Hence usually use the synchronous gauge comoving with CDM: saves one equation (but one more dynamical Einstein equation).

CAMB and CMBFAST use the synchonous gauge comoving with CDM. In CLASS user can choose:

`gauge = synchronous` or `gauge = newtonian`

Equations of motion

Fluids get simple equations of motion because interactions impose unique value of isotropic pressure in each point. Described by $\delta = \delta\rho/\rho$, θ, δp. For perfect fluids $\delta p = w\delta\rho$. By extension we allow for $\delta p = c_s^2 \delta\rho$ with a constant $c_s^2 \neq w$.
For decoupled non-relativistic species, pressure perturbations can be neglected, so effectively equivalent to fluid with $c_s^2 = 0$.

- $\delta' = -(1+w)\theta - 3\frac{a'}{a}\left(c_s^2 - w\right)\delta +$ `metric_continuity` (continuity)
- $\theta' = -\frac{a'}{a}(1-3w)\theta - \frac{w'}{1+w}\theta + \frac{k^2 c_s^2}{1+w}\delta +$ `metric_euler` (Euler)

Equations of motion

For streaming particles: phase space, $f = \bar{f}(\tau, p) + \delta f(\tau, p, \vec{x}, \hat{n})$.

Get δf evolution from Boltzmann equation.

Initially in thermal equilibirum: δf related to $\Theta(\tau, \vec{x})$, simple Boltzmann equation equivalent to continuity+Euler.

Next:

$$\frac{d\ln(ap)}{d\tau} = -\phi' - \frac{\sqrt{p^2+m^2}}{p}\hat{n}\cdot\vec{\nabla}\psi.$$

- ultra-relativistic: temperature still works but $\Theta(\tau, \vec{x}, \hat{n})$ (hierarchy in multipole space, F_l).
- non-relativistic: use $f = \bar{f}\,(1 + \Psi(\tau, p, \vec{x}, \hat{n}))$, Boltzmann in momentum space.

Equations of motion

Einstein equations: 4 scalar equations, but redundant with equations of motion of species (Bianchi identity).

- Newtonian gauge: can be used as constraint equation.
- synchronous gauge comoving with CDM: cannot eliminate differential operator. One dynamical equation, one constraint equation.

In total, same number of equation. However, for high precision in Newtonian gauge, better to replace one constraint equation by one dynamical equation.

Whole ODE: fluid equations, Boltzmann hierarchies, one Einstein equation.

Initial conditions

Different choices (adiabatic, baryon isocurvature, CDM isocurvature, etc..) will be reviewed tomorrow. In each of these cases, all perturbations relate to a single quantity at initial time:

$$\forall i, \qquad A_i(\tau_{\rm ini}, \vec{k}) = a_i \mathcal{R}(\tau_{\rm ini}, \vec{k}).$$

At later time, linearity + isotropy impose

$$\forall i, \qquad A_i(\tau, \vec{k}) = A_i(\tau, k) \mathcal{R}(\tau_{\rm ini}, \vec{k}).$$

Here $A_i(\tau, k)$ is the transfer function normalized to $\mathcal{R}$. Given by the solution of ODE normalized to $\mathcal{R}(\tau_{\rm ini}, \vec{k}) = 1$. Power spectra:

$$\left\langle \left| A_i(\tau, \vec{k}) \right|^2 \right\rangle = (A_i(\tau, k))^2 \left\langle \left| \mathcal{R}(\tau_{\rm ini}, \vec{k}) \right|^2 \right\rangle .$$

Separate tasks for the `perturbation` module, `primordial` module and `spectra` module.

Line-of-sight integral

Boltzmann hierarchy for photons or massless neutrinos: $\left\{\begin{array}{ccc} F_0' & = & ... \\ F_1' & = & ... \\ & ... & \\ F_l' & = & ... \end{array}\right.$

(Ma & Bertschinger: $F_l = 4\Theta_l$)

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Power spectrum in multipole space: $C_l^{TT} = 4\pi \int \frac{dk}{k} \left(\Theta_l(\tau_0, k)\right)^2 \mathcal{P}(k)$. A priori, we need to integrate the Boltzmann hierarchy up to high$l$'s.

Line-of-sight integral

Boltzmann hierarchy for photons or massless neutrinos: $\left\{ \begin{array}{lll} F_0' & = & \dots \\ F_1' & = & \dots \\ & \dots & \\ F_l' & = & \dots \end{array} \right.$

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But clever integration by part of Boltzmann equation gives

$$\Theta_l(\tau_0, k) = \int_{\tau_{\rm ini}}^{\tau_0} d\tau \; S_T(\tau, k) \; j_l(k(\tau_0 - \tau))$$

$$S_T(\tau, k) \equiv \underbrace{g\,(\Theta_0 + \psi)}_{\rm SW} + \underbrace{\left(g\, k^{-2}\theta_{\rm b}\right)'}_{\rm Doppler} + \underbrace{e^{-\kappa}(\phi' + \psi')}_{\rm ISW} + {\rm polarisation}\ .$$

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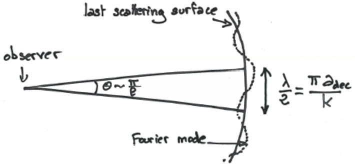
Role of Bessel: projection from Fourier to harmonic space



$$\theta \, d_a(z_{\rm rec}) = \frac{\lambda}{2} \text{ gives precisely } l = k(\tau_0 - \tau_{\rm rec})$$

Instead of calculating all $\Theta_l(\tau, k)$ and store them, we only need to calculate $\Theta_{l\leq 10}(\tau, k)$ and store $S_T(k, \tau)$.
However this source is a non-trivial function (ψ' and polarisation terms)...

Line-of-sight integral

Full source function in CAMB:

```
!Maple fortran output - see scal_eqs.map
        ISW = (4.D0/3.D0*k*EV%Kf(1)*sigma+(-2.D0/3.D0*sigma
            -2.D0/3.D0*etak/adotoa)*k &
              -diff_rhopi/k**2-1.D0/adotoa*dgrho/3.D0+(3.D0*
                  gpres+5.D0*grho)*sigma/k/3.D0 &
              -2.D0/k*adotoa/EV%Kf(1)*etak)*expmmu(j)
!The rest, note y(9)->octg, yprime(9)->octgprime (octopoles)
    sources(1)= ISW +  ((-9.D0/160.D0*pig-27.D0/80.D0*ypol
        (2))/k**2*opac(j)+(11.D0/10.D0*sigma- &
    3.D0/8.D0*EV%Kf(2)*ypol(3)+vb-9.D0/80.D0*EV%Kf(2)*octg
        +3.D0/40.D0*qg)/k-(- &
    180.D0*ypolprime(2)-30.D0*pigdot)/k**2/160.D0)*dvis(j)
        +(-(9.D0*pigdot+ &
    54.D0*ypolprime(2))/k**2*opac(j)/160.D0+pig/16.D0+clxg
        /4.D0+3.D0/8.D0*ypol(2)+(- &
    21.D0/5.D0*adotoa*sigma-3.D0/8.D0*EV%Kf(2)*ypolprime(3)+
        vbdot+3.D0/40.D0*qgdot- &
    9.D0/80.D0*EV%Kf(2)*octgprime)/k+(-9.D0/160.D0*dopac(j)*
        pig-21.D0/10.D0*dgpi-27.D0/ &
    80.D0*dopac(j)*ypol(2))/k**2)*vis(j)+(3.D0/16.D0*ddvis(j
        )*pig+9.D0/ &
    8.D0*ddvis(j)*ypol(2))/k**2+21.D0/10.D0/k/EV%Kf(1)*vis(j
        )*etak
```

Line-of-sight integral

$S_T(k,\tau)$ comes from integration by part of

$$\Theta_l(\tau_0,k) = \int_{\tau_{\rm ini}}^{\tau_0} d\tau \; \Big\{ S_T^0(\tau,k)\; j_l(k(\tau_0-\tau))$$
$$+S_T^1(\tau,k)\; \frac{dj_l}{dx}(k(\tau_0-\tau))$$
$$+S_T^2(\tau,k)\; \frac{1}{2}\left[3\frac{d^2 j_l}{dx^2}(k(\tau_0-\tau)) + j_l(k(\tau_0-\tau))\right]\Big\}$$

CLASS v2.0 stores separately $S_T^1(\tau,k)$, $S_T^2(\tau,k)$, $S_T^2(\tau,k)$, and the transfer module will convolve them individually with respective bessel functions.

$$S_T^0 = g\left(\frac{\delta_g}{4}+\psi\right) + e^{-\kappa}(\phi'+\psi') \qquad S_T^1 = g\frac{\theta_b}{k} \qquad S_T^2 = \frac{g}{8}\,(G_0+G_2+F_2)$$

or

$$S_T^0 = g\left(\frac{\delta_g}{4}+\phi\right) + e^{-\kappa}2\phi' + g'\theta_b + g\theta_b' \qquad S_T^1 = e^{-\kappa}k(\psi-\phi) \qquad S_T^2 = \frac{g}{8}\,(G_0+G_2+F_2)$$

Tensor perturbations

- tensor perturbations present in non-diagonal spatial part of the metric $\delta g_{\mu\nu}$ (2 d.o.f. of GW) and of $\delta T_{\mu\nu}$ of streaming component (decoupled neutrinos and photons).
- no gauge ambiguity for tensors.
- Einstein equations are of course dynamical for GW. For each of the two polarisation states,

$$h'' + 2\frac{a'}{a}h' + k^2 h = f[F_l^\gamma, F_l^\nu] \; .$$

- need to integrate also Boltzmann hierarchies for photons and neutrinos, sourced by GW fields.
- there is also a line-of-sight approach for tensors,

$$\Theta_l(\tau_0, k) = \int_{\tau_{\rm ini}}^{\tau_0} d\tau \; S_T(\tau, k) \; \sqrt{\frac{3(l+2)!}{8(l-2)!}} \frac{j_l(k(\tau_0 - \tau))}{(k(\tau_0 - \tau))^2}$$

$$S_T(\tau, k) = -g\left(\frac{F_{\gamma 0}^{(2)}}{10} + \frac{F_{\gamma 2}^{(2)}}{7} + \frac{3F_4^{(2)}}{70} - \frac{3G_{\gamma 0}^{(2)}}{5} + \frac{6G_{\gamma 2}^{(2)}}{7} - \frac{3G_{\gamma 4}^{(2)}}{70}\right) - e^{-\kappa} h' \; .$$

Polarisation

- Created by Thomson scattering when photons are NOT in thermal equilibrium anymore (recombination, reionisation), such that their "environnement" has a quadrupolar component.
- E-mode and B-mode can be derived from a single source $S_P(\tau, k)$, convolved with different *radial functions* in transfer module.
- absence of any integrated-Sachs-Wolfe-like term:

$$S_P(\tau, k) = \sqrt{6}\, g(\tau)\ P(\tau, k)$$

For scalars:

$$P = \frac{1}{8}\left[F^{(0)}_{\gamma 2} + G^{(0)}_{\gamma 0} + G^{(0)}_{\gamma 2}\right]$$

For tensors:

$$P = -\frac{1}{\sqrt{6}}\left[\frac{F^{(2)}_{\gamma 0}}{10} + \frac{F^{(2)}_{\gamma 2}}{7} + \frac{3F^{(2)}_{4}}{70} - \frac{3G^{(2)}_{\gamma 0}}{5} + \frac{6G^{(2)}_{\gamma 2}}{7} - \frac{3G^{(2)}_{\gamma 4}}{70}\right]$$

T. Tram & JL, JCAP 1310 (2013) 002 [arXiv:1305.3261]

The perturbation module

`source/perturbations.c`

Which sources?

- sources for CMB temperature (decomposed in 1 to 3 terms)
- sources for CMB polarisation (only 1 term)
- sum of metric perturbations $\phi + \psi$ (useful for lensing)
- density perturbations of total non-relativistic matter δ_m (for $P(k)$, NC)
- ϕ, ψ, ϕ', total velocity of non-relativistic matter θ_m (for NC)
- density perturbations of all components with non-zero density $\{\delta_i\}$
- velocity perturbations of all components with non-zero density $\{\theta_i\}$

These different type of sources are called perturbation types and associated to indices

- `index_tp_t0`, `index_tp_t1`, `index_tp_t2`,
- `index_tp_p`,
- `index_tp_phi_plus_psi`, `index_tp_phi`, `index_tp_psi`, `index_tp_phi_prime`
- `index_tp_delta_m`, `index_tp_theta_m`
- `index_tp_delta_g`, `index_tp_delta_cdm`, etc.
- `index_tp_thelta_g`, `index_tp_thelta_cdm`, etc.

of size `tp_size`.

Which sources?

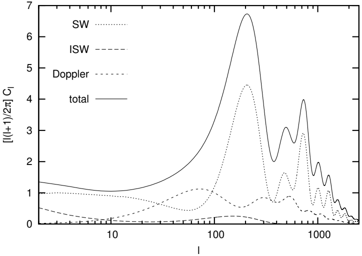
Computing/storing each of these sources is decided automatically by the code depending on what the user asked:

- which output? CMB temperature, CMB polarisation, CMB lensing, matter power spectrum, all densities $\{\delta_i\}$, all velocities $\{\theta_i\}$, C_l's of number count in redshift bin, C_l's of cosmic shear in redshift bin, corresponding respectively to `output = tCl, pCl, lCl, mPk, dTk, vTk, nCl, sCl`
- which modes? scalars, vectors, tensors, associated to indices `index_md_scalars`, `index_md_vectors`, `index_md_tensors`.
- Different number of sources for each mode, `tp_size[index_mode]`
- The code must consider $S(k, \tau)$ for each type, each mode, but also each initial condition (adiabatic, CDM isocurvature, neutrino isocurvature, etc., again associated to indices).
- could define
  `ppt->sources[index_md][index_ic][index_type][index_tau][index_k]`
  but better to use
  `ppt->sources[index_md][index_ic * ppt->tp_size[index_md] + index_type][index_tau * ppt->k_size + index_k]`

Which sources?

Possibility to take into account only a few contribution to the sources:

- For CMB temperature `tCl` (to understand physically different effects):
  `temperature contributions = tsw, eisw, lisw, dop, pol`
  `early/late isw redshift = 50`



- For number count `nCl` (to understand physically different effects, and also to speed up, keeping only leading contributions, see `arxiv:1307.1459` or Bonvin & Durrer `arxiv:1105.5280`):
  `number count contributions = density, rsd, lensing, gr`

Main functions in perturbations.c

Very few external functions:

- `perturb_sources_at_tau()`, actually never called because the subsequent modules read `ppt->sources[...]` without interpolating
- `perturb_init()`, which ultimate goals is to fill `ppt->sources[index_md][index_ic * ppt->tp_size[index_md] + index_type][index_tau * ppt->k_size + index_k]`
- `perturb_free()`, which frees the memory allocated in the structure `ppt`.

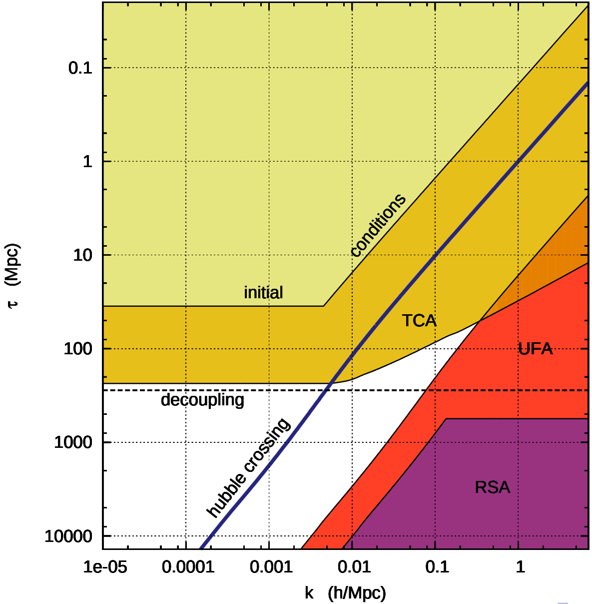
Main tasks in perturb_init()

- define all indices with `perturb_indices()`
- find an optimal time-sampling of the sources, based on the variation rate of background and thermodynamical quantitites, and on k.
- loop over all modes, initial conditions, wavenumbers
- for each of them, call `perturb_solve()` to compute $S(k,\tau)$ of each type

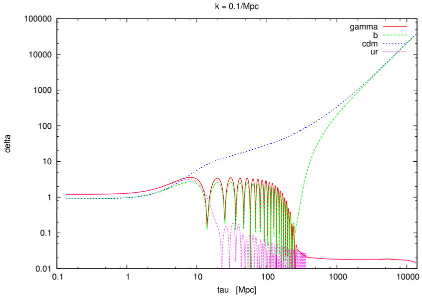
Main tasks in perturb_solved()

- to find when approximation schemes (tight coupling, ultra-relativistic fluid approximation, radiation streaming approximation) must be switched off or switched on.
- inside each interval where the approximation does not change, to integrate the system of cosmological perturbations `dy[index_pt] = ... y[index_pt]`, using `perturb_initial_conditions()`, `perturb_derivs()`, `perturb_total_stress_energy()`, `perturb_einstein()`
- when the approximation scheme changes, manage to redefine `y[index_pt]` and ensure continuity
- each times that we cross a value of τ where we wish to sample the sources, compute them and store them with `perturbations_sources()`

Approximation schemes



Printing the perturbation evolution



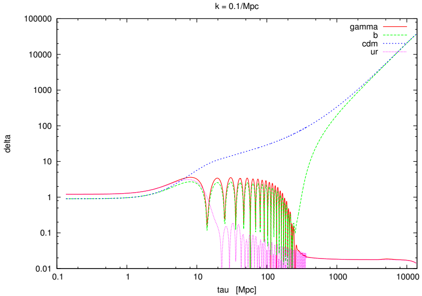
Execute e.g. `./class myinput.ini` including in the input file:

```
k_output_values = 0.01,0.1,1
root = output/toto_
```

If scalars are requested: the perturbation module will write files `output/toto_perturbations_k0_s.dat`, ..., with an explicit header:

```
#scalar perturbations for mode k = 1.001486417109e-02 Mpc^(-1)
# tau [Mpc] a delta_g theta_g shear_g pol0_g pol1_g pol2_g delta_b
theta_b psi phi delta_ur theta_ur shear_ur delta_cdm theta_cdm
```

Printing the perturbation evolution



Execute e.g. `./class myinput.ini` including in the input file:

```
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```

If tensors are requested: similar output written in files `output/toto_perturbations_k0_t.dat`, ...

Plotting a source with test/test_perturbations

```
input_init_from_arguments(argc, argv,&pr,&ba,&th,&pt,&tr,&pm
    ,&sp,&nl,&le,&op,errmsg);

background_init(&pr,&ba);
thermodynamics_init(&pr,&ba,&th);
perturb_init(&pr,&ba,&th,&pt);

/* choose a mode (scalar, tensor, ...) */
int index_mode=pt.index_md_scalars;

/* choose a type (temperature, polarization, grav. pot.,
    ...) */
int index_type=pt.index_tp_t0;

/* choose an initial condition (ad, bi, cdi, nid, niv, ...)
    */
int index_ic=pt.index_ic_ad;
```

Plotting a source with test/test_perturbations.c

```
output=fopen("output/source.dat","w");
fprintf(output,"#   k         tau         S\n");

for (index_k=0; index_k < pt.k_size; index_k++) {
  for (index_tau=0; index_tau < pt.tau_size; index_tau++) {

    fprintf(output,"%e %e %e\n",
      pt.k[index_k],
      pt.tau_sampling[index_tau],
      pt.sources[index_mode]
      [index_ic * pt.tp_size[index_mode] + index_type]
      [index_tau * pt.k_size + index_k]
      );
  }
}
```

For instance you can type:

```
> make test_perturbations
> ./test_perturbations my_input.ini
```

Exercises

Exercise III

Compare the evolution of $\phi(k,\tau)$ and $\psi(k,\tau)$ for $k = 0.01, 0.1/\mathrm{Mpc}$. Check that they are not equal on super-Hubble scales. To understand why, plot $a^2(\bar{\rho}_\nu + \bar{p}_\nu)\sigma_\nu$ versus time and check that the results are consistent with the Einstein equation
$k^2(\phi - \psi) = 12\pi G a^2(\bar{\rho} + \bar{p})\sigma_{\mathrm{tot}}$.

Physics of acoustic oscillations

From equations to shape of CMB temperature spectrum.
(References: *TASI lecture notes* arXiv:1302.4640 , *Neutrino Cosmology* book section 5.1)

Tight-coupling limit:

$$\Theta_0'' + \frac{R'}{1+R}\Theta_0' + k^2 c_\mathrm{s}^2 \Theta_0 = -\frac{k^2}{3}\psi + \frac{R'}{1+R}\phi' + \phi'' \ .$$

with

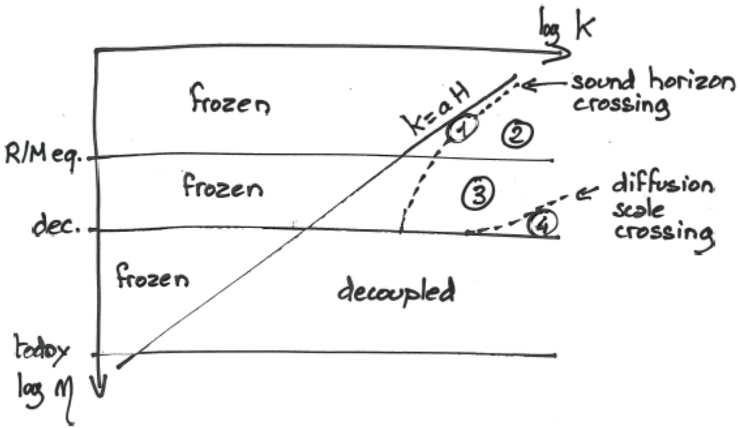
$$c_\mathrm{s}^2 = \frac{1}{3(1+R)} \ , \quad R \equiv \frac{4\bar{\rho}_\mathrm{b}}{3\bar{\rho}_\gamma} \propto a \ .$$

When R and metric derivatives are constant, acoustic oscillations $\propto \cos(k c_s \tau)$ around zero-point

$$\Theta_0^{zero-point} = -(1+R)\psi.$$

Hubble crossing: $\lambda \sim 1/H$, i.e. $k \sim aH$, i.e. $k\tau \sim 1$.

Physics of acoustic oscillations

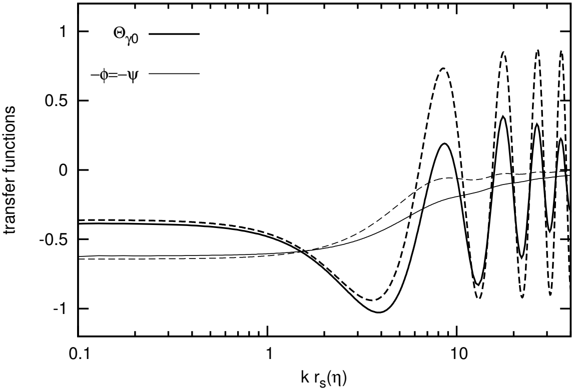


Initially static with $\Theta_0 = -\frac{2}{3}\psi$. Then:

1. start to oscillate around $\Theta_0^{zero-point} = -(1+R)\psi$.
2. metric fluctuations decay, boosting amplitude, then stationary oscillations with $\Theta_0^{zero-point} = 0$.
3. baryon damping and decay of sound speed: slow damping of oscillations.
4. diffusion erases perturbations exponentially.

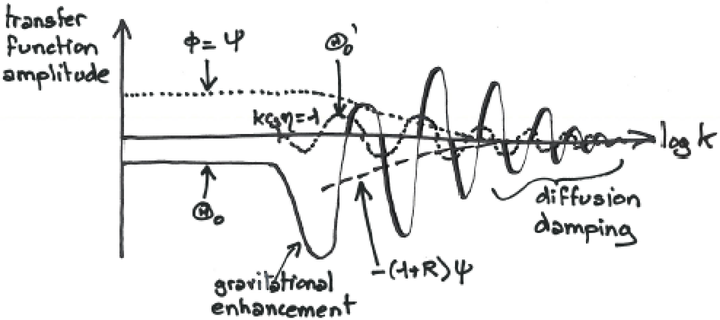
Physics of acoustic oscillations

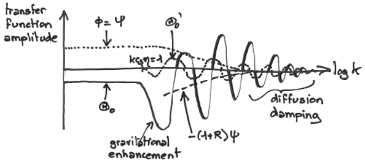
Snapshot of perturbations at given time (here, equality and decoupling)

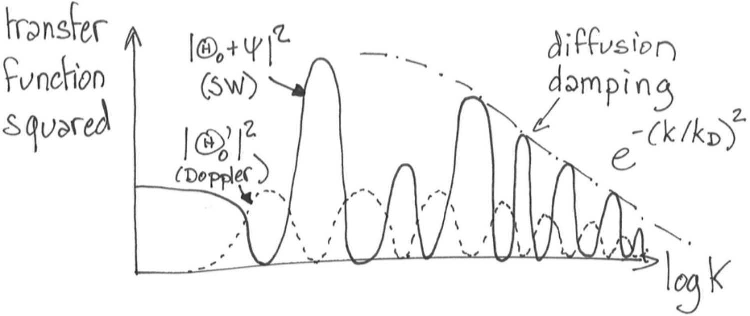


From *Neutrino cosmology* book, see later how to get snapshots.

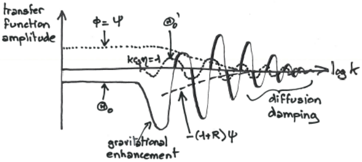
Snapshot of perturbations at equality:

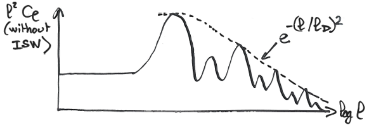


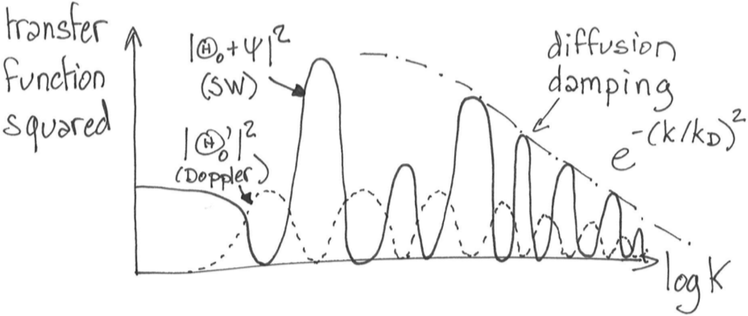

transfer function amplitude
$\Phi = \Psi$
Θ_0'
$k c_s \eta = 1$
log k
Θ_0
diffusion damping
$-(1+R)\Psi$
gravitational enhancement


transfer function squared
$|\Theta_0 + \Psi|^2$
(SW)
$|\Theta_0'|^2$
(Doppler)
diffusion damping
$e^{-(k/k_D)^2}$
log K

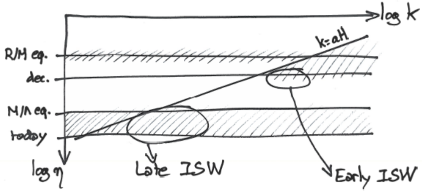
Physics of acoustic oscillations

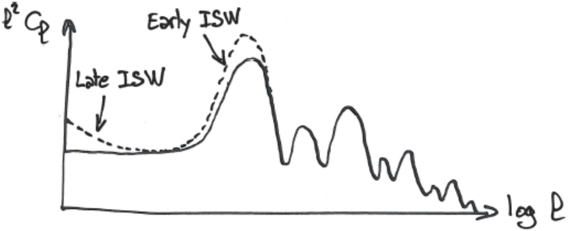


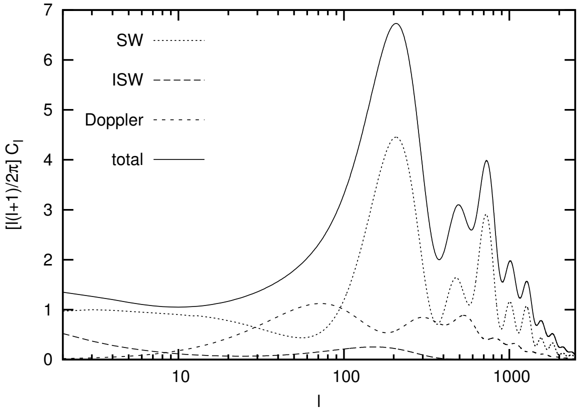




Physics of acoustic oscillations





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temperature contributions = tsw, eisw, lisw, dop, pol
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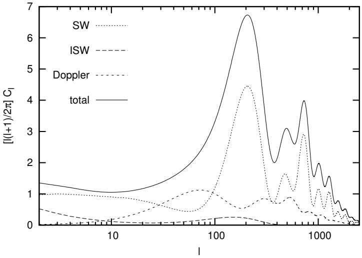
In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



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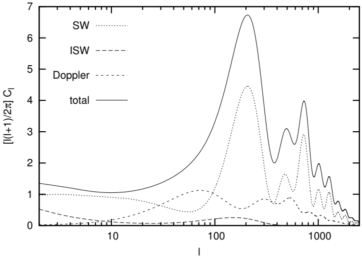

- **(C1)** Peak location: depends on angle $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm rec})$

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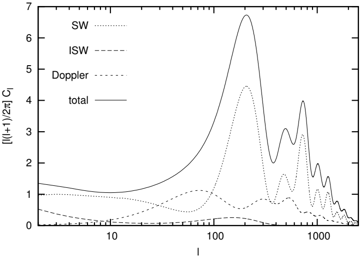

- **(C1)** Peak location: depends on angle $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm rec})$
- **(C2)** Ratio of first-to-second peak: gravity-pressure balance in fluid, ω_b/ω_γ

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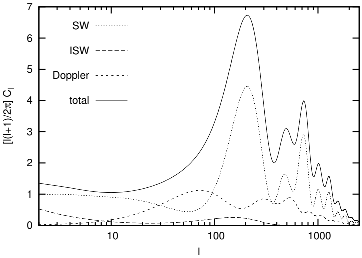
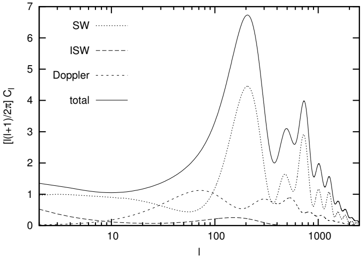


- **(C1)** Peak location: depends on angle $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm rec})$
- **(C2)** Ratio of first-to-second peak: gravity-pressure balance in fluid, ω_b/ω_γ
- **(C3)** Time of equality: amplitude of first peaks (no boosting during MD), effect enhanced for 1st peak (early ISW); depends on $a_{\rm dec}/a_{\rm eq}$

In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



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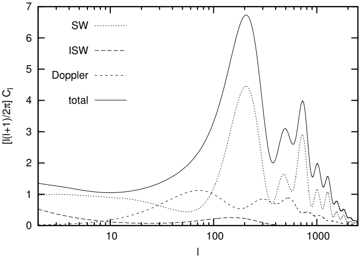
- **(C4)** Enveloppe of high-l peaks: diffusion scale, angle $\theta = \lambda_d(\tau_{\rm dec})/d_A(\tau_{\rm dec})$

In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



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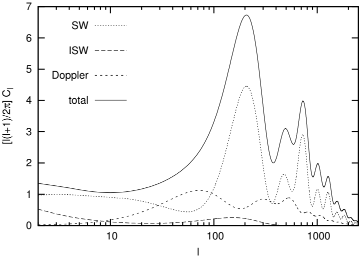
- **(C5)** Global amplitude: A_s

In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



- **(C1)** Peak location: depends on angle $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm rec})$
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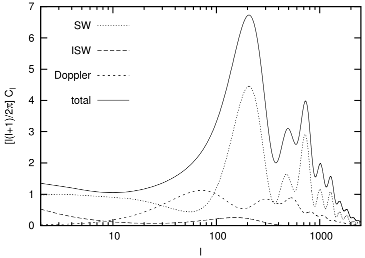
- **(C5)** Global amplitude: A_s
- **(C6)** Global tilt: n_s

In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



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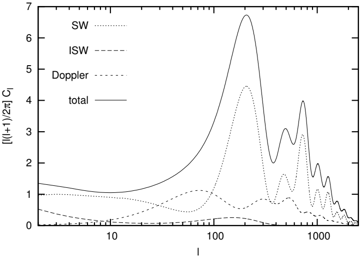
- **(C4)** Enveloppe of high-l peaks: diffusion scale, angle $\theta = \lambda_d(\tau_{\rm dec})/d_A(\tau_{\rm dec})$
- **(C5)** Global amplitude: A_s
- **(C6)** Global tilt: n_s
- **(C7)** Slope of Sachs-Wolfe plateau (beyond tilt effect): late ISW, z_Λ

In neutrinoless ΛCDM model, C_l^{TT} controlled by 8 effects/quantitites:



- **(C1)** Peak location: depends on angle $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm rec})$
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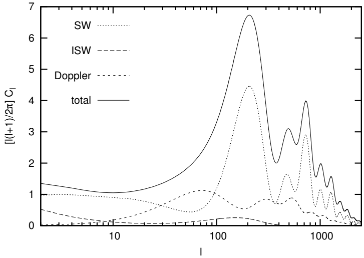
- **(C3)** Time of equality: amplitude of first peaks (no boosting during MD), effect enhanced for 1st peak (early ISW); depends on $a_{\rm dec}/a_{\rm eq}$
- **(C4)** Enveloppe of high-l peaks: diffusion scale, angle $\theta = \lambda_d(\tau_{\rm dec})/d_A(\tau_{\rm dec})$
- **(C5)** Global amplitude: A_s
- **(C6)** Global tilt: n_s
- **(C7)** Slope of Sachs-Wolfe plateau (beyond tilt effect): late ISW, z_Λ
- **(C8)** Relative amplitude for $l \gg 40$ w.r.t $l \ll 40$: optical depth τ_{reio}

In terms of parameters $\{\omega_m, \omega_b, \Omega_\Lambda, A_s, n_s, \tau_{reio}\}$:

(with $h = \sqrt{\omega_m/(1-\Omega_\Lambda)}$ and $\omega_m = \omega_b + \omega_c$)



- **(C1)** Peak location: $\theta = d_s(\tau_{\rm dec})/d_A(\tau_{\rm dec})$ — ω_m, ω_b, Ω_Λ
- **(C2)** Ratio of first-to-second peak: ω_b/ω_γ — ω_b
- **(C3)** Time of equality: $z_{eq} = \omega_m/\omega_\gamma$ — ω_m
- **(C4)** Enveloppe of high-l peaks: $\theta = \lambda_d(\tau_{\rm dec})/d_A(\tau_{\rm dec})$ — ω_m, ω_b, Ω_Λ
- **(C5)** Global amplitude: A_s — A_s
- **(C6)** Global tilt: n_s — n_s
- **(C7)** Slope of Sachs-Wolfe plateau: z_Λ — Ω_Λ
- **(C8)** Relative amplitude for $l \gg 40$ w.r.t $l \ll 40$: optical depth — τ_{reio}

Exercise IV

There exist several ways to parametrise modifications of gravity. For instance, people often study the effect of a function $\mu(k,\tau)$ inserted in the Poisson equation, giving in the synchronous gauge:

$$k^2\eta - \frac{1}{2}\frac{a'}{a}h' = \mu(k,\tau)\ 4\pi G a^2 \bar{\rho}_{\rm tot}\delta_{\rm tot}\ .$$

Localise the above equation and implement, for instance, $\mu = 1 + a^3$. Print the evolution of ϕ and ψ in the standard and modified models, and conclude that the C_l^{TT}'s should be affected only through the late ISW effect. Get a confirmation by comparing directly the C_l's.